



Radar-Enabled Recovery of the Sutter's Mill Meteorite, a Carbonaceous Chondrite Regolith Breccia

Peter Jenniskens *et al.*
Science **338**, 1583 (2012);
DOI: 10.1126/science.1227163

This copy is for your personal, non-commercial use only.

If you wish to distribute this article to others, you can order high-quality copies for your colleagues, clients, or customers by [clicking here](#).

Permission to republish or repurpose articles or portions of articles can be obtained by following the guidelines [here](#).

The following resources related to this article are available online at www.sciencemag.org (this information is current as of December 20, 2012):

Updated information and services, including high-resolution figures, can be found in the online version of this article at:

<http://www.sciencemag.org/content/338/6114/1583.full.html>

Supporting Online Material can be found at:

<http://www.sciencemag.org/content/suppl/2012/12/19/338.6114.1583.DC1.html>

A list of selected additional articles on the Science Web sites **related to this article** can be found at:

<http://www.sciencemag.org/content/338/6114/1583.full.html#related>

This article **cites 86 articles**, 10 of which can be accessed free:

<http://www.sciencemag.org/content/338/6114/1583.full.html#ref-list-1>

This article appears in the following **subject collections**:

Planetary Science

http://www.sciencemag.org/cgi/collection/planet_sci

Radar-Enabled Recovery of the Sutter's Mill Meteorite, a Carbonaceous Chondrite Regolith Breccia

Peter Jenniskens,^{1,2*} Marc D. Fries,³ Qing-Zhu Yin,⁴ Michael Zolensky,⁵ Alexander N. Krot,⁶ Scott A. Sandford,² Derek Sears,² Robert Beauford,⁷ Denton S. Ebel,⁸ Jon M. Friedrich,^{8,9} Kazuhide Nagashima,⁶ Josh Wimpenny,⁴ Akane Yamakawa,⁴ Kunihiko Nishiizumi,¹⁰ Yasunori Hamajima,¹¹ Marc W. Caffee,¹² Kees C. Welten,¹⁰ Matthias Laubenstein,¹³ Andrew M. Davis,^{14,15} Steven B. Simon,¹⁴ Philipp R. Heck,^{14,15} Edward D. Young,¹⁶ Issaku E. Kohl,¹⁷ Mark H. Thiemens,¹⁸ Morgan H. Nunn,¹⁸ Takashi Mikouchi,¹⁹ Kenji Hagiya,²⁰ Kazumasa Ohsumi,²¹ Thomas A. Cahill,²² Jonathan A. Lawton,²² David Barnes,²² Andrew Steele,²³ Pierre Rochette,²⁴ Kenneth L. Verosub,⁴ Jérôme Gattacceca,²⁴ George Cooper,² Daniel P. Glavin,²⁵ Aaron S. Burton,^{25,26} Jason P. Dworkin,²⁵ Jamie E. Elsila,²⁵ Sandra Pizzarello,²⁷ Ryan Ogliore,⁶ Phillipe Schmitt-Kopplin,^{28,29} Mourad Harir,²⁸ Norbert Hertkorn,²⁸ Alexander Verchovsky,³⁰ Monica Grady,³⁰ Keisuke Nagao,³¹ Ryuji Okazaki,³² Hiroyuki Takechi,³² Takahiro Hiroi,³³ Ken Smith,³⁴ Elizabeth A. Silber,³⁵ Peter G. Brown,³⁵ Jim Albers,¹ Doug Klotz,³⁶ Mike Hankey,³⁷ Robert Matson,³⁸ Jeffrey A. Fries,³⁹ Richard J. Walker,⁴⁰ Igor Puchtel,⁴⁰ Cin-Ty A. Lee,⁴¹ Monica E. Erdman,⁴¹ Gary R. Eppich,⁴² Sarah Roeske,⁴ Zelimir Gabelica,⁴³ Michael Lerche,⁴⁴ Michel Nuevo,^{1,2} Beverly Girten,² Simon P. Worden² (the Sutter's Mill Meteorite Consortium)

Doppler weather radar imaging enabled the rapid recovery of the Sutter's Mill meteorite after a rare 4-kiloton of TNT-equivalent asteroid impact over the foothills of the Sierra Nevada in northern California. The recovered meteorites survived a record high-speed entry of 28.6 kilometers per second from an orbit close to that of Jupiter-family comets (Tisserand's parameter = 2.8 ± 0.3). Sutter's Mill is a regolith breccia composed of CM (Mighei)-type carbonaceous chondrite and highly reduced xenolithic materials. It exhibits considerable diversity of mineralogy, petrography, and isotope and organic chemistry, resulting from a complex formation history of the parent body surface. That diversity is quickly masked by alteration once in the terrestrial environment but will need to be considered when samples returned by missions to C-class asteroids are interpreted.

On 22 April 2012, the KBBX (Beale Air Force Base, California), KDAX (Sacramento, California), and KRGX (Reno, Nevada) weather radars of the U.S. National Climatic Data Center's NEXRAD network (*J*) detected radial Doppler shifts in four sweeps, following a fast-moving daytime fireball seen over much of California and Nevada at 14:51:12 to 17 UTC (Fig. 1). The falling meteorites were

identified from a downward sequence of subsequent detections, small-scale turbulence, and widely variable spectrum width values correlated in time, location, and direction with eyewitness reports of the fireball.

Under the radar footprint over the townships of Coloma and Lotus in El Dorado County, California, the first three pieces of the meteorite were recovered on 24 April, before heavy rain hit

the area (2). One meteorite fell at Sutter's Mill (SM), the gold discovery site that initiated the California Gold Rush. Two months after the fall, SM find numbers were assigned to the 77 meteorites listed in table S3 (3), with a total mass of 943 g. The biggest meteorite is 205 g.

This is a tiny fraction of the pre-atmospheric mass, based on the kinetic energy derived from infrasound records. Eyewitnesses reported hearing a loud boom followed by a deep rumble. Infrasound signals (table S2A) at stations I57US and I56US of the International Monitoring System (4), located ~770 and ~1080 km from the source, are consistent with stratospherically ducted arrivals (5). The combined average periods of all phase-aligned stacked waveforms at each station of 7.6 s correspond to a mean source energy of 4.0 (−2.2/+3.4) kT of TNT, using the multistation period yield relation from (5). This was the most energetic reported bolide falling on land globally (6) since the 1.2-kT impact of asteroid 2008 TC₃ over Sudan in 2008 (7).

Seismic data suggest a point source altitude of 54.8 ± 10.9 km above mean sea level, estimated from impulsive phase arrivals of the air blast at eight seismograph stations (8) by applying a simple half-space sonic velocity model and direct ray paths to a standard earthquake location code (table S2B).

This altitude corresponds to a persistent flare detected in a set of three photographs from Rancho Haven, north of Reno, Nevada (fig. S1). Triangulation with videos from Johnsondale, California, and Incline Village, Nevada (figs. S2 and S3), shows that the bolide was first detected at 90 km approaching from the east, had a broad peak in brightness around 56 km, and detonated at 47.6 ± 0.7 km (Table 1). Even in the daytime sky, a great many fragments were detected down to 30 km.

The entry speed is twice that of typical triangulated falls from which meteorites have been recovered (table S1). SM has the highest disruption altitude on record. With an entry velocity of 28.6 km/s, the infrasound-derived kinetic energy corresponds to a pre-atmospheric mass of ~40,000

¹SETI Institute, 189 Bernardo Avenue, Mountain View, CA 94043, USA. ²NASA Ames Research Center, Moffett Field, CA 94035, USA. ³Planetary Science Institute, Tucson, AZ 85719–2395, USA. ⁴Department of Geology, University of California at Davis, Davis, CA 95616, USA. ⁵Astromaterials Research and Exploration Science, NASA Johnson Space Center, Houston, TX 77058, USA. ⁶Hawai'i Institute of Geophysics and Planetology and Astrobiology Institute, University of Hawai'i at Mānoa, Honolulu, HI 96822, USA. ⁷Arkansas Center for Space and Planetary Sciences, University of Arkansas, AR 72701, USA. ⁸Department of Earth and Planetary Sciences, American Museum of Natural History, New York, NY 10024, USA. ⁹Department of Chemistry, Fordham University, Bronx, NY 10458, USA. ¹⁰Space Sciences Laboratory, University of California, Berkeley, CA 94720–7450, USA. ¹¹Low Level Radioactivity Laboratory, Kanazawa University, Nomi, Ishikawa 923-1224, Japan. ¹²Department of Physics, Purdue University, West Lafayette, IN 47907, USA. ¹³Istituto Nazionale di Fisica Nucleare, Laboratori Nazionali del Gran Sasso, I-67100 Assergi, Italy. ¹⁴Department of the Geophysical Sciences, Enrico Fermi Institute and Chicago Center for Cosmochemistry, The University of Chicago, Chicago, IL 60637, USA. ¹⁵Robert A. Pritzker Center for Meteoritics and

Polar Studies, Field Museum of Natural History, Chicago, IL 60605, USA. ¹⁶Department of Earth and Space Sciences, University of California Los Angeles, Los Angeles, CA 90095–1567, USA. ¹⁷Jet Propulsion Laboratory/California Institute of Technology, Pasadena, CA 91109, USA. ¹⁸Department of Chemistry and Biochemistry, University of California San Diego, La Jolla, CA 92093, USA. ¹⁹Department of Earth and Planetary Science, University of Tokyo, Hongo, Bunkyo-ku, Tokyo 113-0033, Japan. ²⁰Graduate School of Life Science, University Hyogo, Hyogo 678-1297, Japan. ²¹Japan Synchrotron Radiation Research Institute, Sayo-cho, Hyogo 679-5189, Japan. ²²Department of Physics, University of California at Davis, Davis, CA 95616, USA. ²³Geophysics Laboratory, Carnegie Institution of Washington, Washington, DC 20015, USA. ²⁴Centre Européen de Recherche et d'Enseignement, CNRS/Aix-Marseille Université, F-13545 Aix-en-Provence, France. ²⁵NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA. ²⁶Oak Ridge Associated Universities, Greenbelt, MD 20771, USA. ²⁷Arizona State University, Tempe, AZ 85287, USA. ²⁸Helmholtz Zentrum München, D-85764 München, Germany. ²⁹Analytische Lebensmittel Chemie, Technische Universität München, Freising, Germany. ³⁰Planetary and Space Sciences Research Institute, Open University, Milton Keynes MK7

6AA, UK. ³¹Geochemical Research Center, University of Tokyo, Bunkyo-ku, Tokyo 113-0033, Japan. ³²Department of Earth and Planetary Sciences, Kyushu University, Hakozaki, Fukuoka 812-8581, Japan. ³³Department of Geological Sciences, Brown University, Providence, RI 02912, USA. ³⁴Nevada Seismological Laboratory, University of Nevada, Reno, NV 89557, USA. ³⁵University of Western Ontario, London, Ontario N6A 3K7, Canada. ³⁶Space Science for Schools, Incline Village, NV 89451, USA. ³⁷American Meteor Society, Geneseo, NY 14454, USA. ³⁸Science Applications International Corporation, Seal Beach, CA 90740, USA. ³⁹U.S. Air Force Weather Agency, 1st Weather Group, Offutt Air Force Base, NE 68113, USA. ⁴⁰Department of Geology, University of Maryland, College Park, MD 20742, USA. ⁴¹Department of Earth Science, Rice University, Houston, TX 77005, USA. ⁴²Glenn Seaborg Institute, Lawrence Livermore National Laboratory, Livermore, CA 94550, USA. ⁴³Université de Haute Alsace, F-68093 Mulhouse, France. ⁴⁴McClellan Nuclear Research Center, University of California at Davis, McClellan, CA 95652, USA.

*To whom correspondence should be addressed. E-mail: petrus.m.jenniskens@nasa.gov

(range 20,000 to 80,000) kg. Counter to intuition, the catastrophic disruption was key to meteorite survival from this fast entry (7). So far, ~ 0.1 kg/km has been recovered along the trend line, for an estimated total fallen mass ≥ 1.7 kg. This is far less than that recovered from the similar-sized but slower-impacting 2008 TC₃: 1.2 kg/km and 39 kg (7). An airship search did not find impact scars from falling kilogram-sized meteorites.

The pre-atmospheric orbit (Table 1) had low inclination and high eccentricity, with aphelion close to the orbit of Jupiter and perihelion (q) approaching the orbit of Mercury (fig. S5). Semimajor axis, eccentricity, and inclination are strikingly similar to the preliminary values reported for the CM-type carbonaceous chondrite Maribo (9) by Haack *et al.* (10). The entry conditions, too, were much alike. Upon request, W. Singer and G. Stober provided the seven Juliusruh radar head-echo range and direction positions of the early 112- to 79-km altitude part of this fall (table S1), which translated into the orbital elements listed in Table 1. SM and Maribo have lower perihelion distance and higher-eccentricity orbits than all other known (mostly ordinary chondrite) falls (table S1).

That distinction may be due to a bias toward a more recent evolution into Earth-crossing orbits for this population. SM has a cosmic ray exposure (CRE) age at the extreme low end of the CM2 chon-

drite CRE age distribution, which as a group is younger than all other classes of meteorites except lunar meteorites (11). An age of 0.10 ± 0.04 million years (My) was obtained from the measured ^{26}Al activity (with a 0.705 million-year half-life) of 3.8 ± 0.8 dpm/kg in SM36 (table S20). He and Ar isotopic ratios in SM43 and SM51 show no clear excess of a cosmogenic noble gas component, but a small excess of cosmogenic Ne is observed as deviations from two mixing lines between the trapped components; solar Ne – P3 Ne and P3 Ne – Ne E (12, 13) (fig. S27). The low concentration of cosmogenic ^{21}Ne [$1.02 \pm 0.11 \times 10^{-10}$ cm³/g at standard temperature and pressure (STP)] indicates a very short CRE age of 0.051 ± 0.006 My (the average of four ages in table S18), if a production rate of 2×10^{-9} cm³/g/My at STP for the near surface of a large object is adopted (14).

The SM and Maribo orbits have a Tisserand's parameter with respect to Jupiter (T_J) that borders those of asteroids ($T_J > 3$) and Jupiter family comets ($T_J = 2$ to 3). Visual observations of Murchison (CM) also point to approach on a low-inclined orbit (15). A possible Jupiter-family comet origin is intriguing (10) [for a review, see (15)], especially because the CM-like micrometeorites are dynamically linked to Jupiter-family comets (16). Finding two asteroids on orbits at $q = 0.47$ astronomical unit (AU) could mean that both are part

of an old ~ 0.1 -My-old meteoroid stream, perhaps related to 2P/Encke (10), but only if CM chondrites survive longer than typical Taurid meteoroids.

Until other evidence of hydrothermal alteration in Jupiter-family comets is found (15), an origin in the asteroid belt is more likely. The asteroid-family source has low inclination and is close to the 3:1 mean motion resonance with Jupiter, the CRE age leaving little time for thermal drag forces to move the semimajor axis of a main-belt asteroid into resonance at 2.5 AU. A suitable candidate is the (495) Eulalia family, recently proposed as a source of near-Earth C-class asteroids (17). Unlike ordinary chondrites ejected from the 3:1 resonance, which tend to collide with Earth on higher-impact-probability perihelion distance $q \sim 1$ AU orbits of lower eccentricity (table S1), the fractured CM chondrites may disintegrate too rapidly to evolve into such orbits. The short exposure age of SM as compared to other CM chondrites could mean it had already broken from a larger precursor while evolving into an SM-like orbit.

The reflectance spectrum of SM12 (fig. S8) is a good match to the Hayabusa 2 mission's target, asteroid 1999 JU₃, over the measured 0.38- to 0.92- μm range when normalized at 0.55 μm (fig. S8). The meteorite's albedo at 0.55 μm is low, 2.5 to 4.0% at the standard 30° incidence and 0° emergence angles. Carbonates are abundant in this sample and produce strong infrared absorption bands near 1450 and 875 cm⁻¹, stronger than those seen in CM2 Murchison.

SM12 has a compression strength of 82 ± 6 MPa, as compared to ~ 30 MPa for other CM2s (18). The impacting asteroid fragmented at a dynamic pressure of only 0.9 MPa at 48 km, presumably because of internal cracks (19). X-ray computed tomography scanning (20) of SM3, 9, 18, 51, 54, and 73 at the 12- to 30- μm /voxel edge showed abundant fractures through stones and abundant fractures within some lithic fragments, terminating at the fragment edges (fig. S28). Volume measurements (table S19) yielded densities of 2.27 ± 0.07 g/cm³, similar to an average CM2 density of 2.20 g/cm³ (range 1.88 to 2.47) (21). He ideal-gas pycnometry of SM19 gave a bulk density of 2.31 ± 0.04 g/cm³, a grain density of 3.34 ± 0.02 g/cm³, and a high porosity of $31.0 \pm 1.4\%$ (22).

This density and the infrasound-derived kinetic energy yield an asteroid diameter of 2.5 to

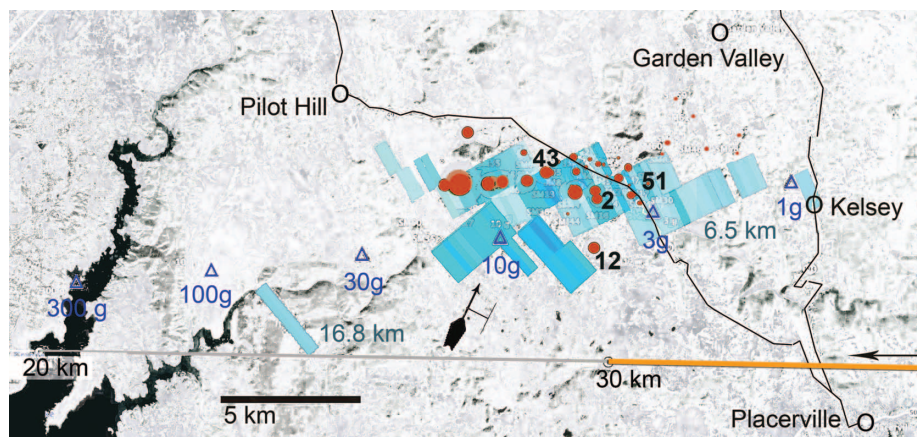


Fig. 1. Map of ground-projected fireball trajectory solution (orange line), radar Doppler reflectivity (light blue; -5 dBZ for pale blue to $+15$ dBZ for deepest blue), and meteorite find locations near Sutter's Mill (red dots are of size proportional to mass; relevant finds are labeled with SM number). Blue triangles are the calculated impact locations for breakup at 48 km and wind drift from 13 m/s winds to azimuth 27° between 18 and 2 km altitude.

Table 1. Atmospheric trajectory and pre-atmospheric orbit for the SM and Maribo CM chondrites. Angular elements are for equinox J2000.0.

Atmospheric trajectory	SM	Maribo	Pre-atmospheric orbit	SM	Maribo
H_b (beginning height - km)	90.2 ± 0.4	111.8 ± 0.4	T_J (Tisserand's parameter)	2.81 ± 0.32	3.04 ± 0.32
H_m (broad maximum - km)	~ 56	~ 58	a (semimajor axis - AU)	2.59 ± 0.35	2.34 ± 0.29
H_f (disruption - km)	47.6 ± 0.7	37.3 ± 0.6	e (eccentricity)	0.824 ± 0.020	0.795 ± 0.026
H_e (end height - km)	~ 30.1	~ 32	q (perihelion distance - AU)	0.456 ± 0.022	0.481 ± 0.010
V_∞ (entry speed - km/s)	28.6 ± 0.6	28.0 ± 0.7	ω (argument of perihelion - °)	77.8 ± 3.2	99.0 ± 1.4
h (entry elevation angle - °)	26.3 ± 0.5	30.2 ± 0.5	Ω (longitude of ascending node - °)	32.77 ± 0.06	117.64 ± 0.05
a_z (entry azimuth angle from south - °)	272.5 ± 0.4	276.2 ± 0.2	i (inclination - °)	2.38 ± 1.16	0.72 ± 0.98
V_g (geocentric entry speed - km/s)	26.0 ± 0.7	25.4 ± 0.8	Q (aphelion distance - AU)	4.7 ± 0.7	4.2 ± 0.6
Ra_g (geocentric right ascension of radiant - °)	24.0 ± 1.3	124.6 ± 1.0	T_p (perihelion time)	2012-03-09.1	2008-12-03.6
Dec_g (geocentric declination of radiant - °)	12.7 ± 1.7	18.8 ± 1.6	Epoch (UT)	2012-04-22.620	2009-01-17.798

4.0 m. Concentrations of the cosmogenic radionuclide ^{60}Co in three fragments—SM18, SM36 (table S20A), and SM43—and model calculations (23) confirm that the asteroid's pre-atmospheric size was >0.9 m. The fireball's peak luminosity (-18 to -20 magnitude) at the first broad maximum suggests a size of 1.8 to 3.5 m.

Textural and compositional variety. Because of its large size, SM provides insight into the variety of materials present at the surface of its parent body. Individual SM meteorites have differing magnetic susceptibilities (χ in 10^{-9} m^3/kg). Ten different stones (table S4) suggest a bimodal distribution clustering around $\log_{10}\chi = 4.03$ and 4.26 . The primary magnetic mineral is magnetite (Fe_3O_4), with concentrations of 2.0 and 3.3 weight % (wt %), respectively. These magnetite concentrations are intermediate between typical CM2s (~ 1 wt %) and magnetite-rich C2s or anomalous CMs (>6 wt %) (24). SM2 has a stable natural remnant magnetization, probably extraterrestrial in origin, corresponding to a magnetic field paleointensity of ~ 3 μT , comparable to other carbonaceous chondrites (25).

SM is a regolith breccia. Like all CM chondrites (26), SM contains Ne from solar wind implantation in a surface regolith (fig. S27). Unlike most other CM chondrites, the brecciated nature of that regolith is evident: SM2, 18, 47, 48, 51, and 54 contain angular to rounded clasts embedded in a fine-grained comminuted matrix seen visually (Fig. 2A), by x-ray and backscattered electron mapping (Fig. 2B and fig. S13), and by x-ray and neutron computed tomography (figs. S28 and S30).

The classification of SM as a CM chondrite is confirmed by whole-rock chemistry (Fig. 3) and by O (Fig. 4 and fig. S21), Os (supplementary text S2.6), and Cr (supplementary text S2.7) isotopic compositions. The Os isotopic compositions and highly siderophile element abundances are well within the range of CM chondrites (figs. S16 and S17). However, the Re-Os isotopic systematics indicate minor, what was probably recent, open-system behavior of these two elements, as seen in other chondrites (27). SM plots in the field of CM chondrites (28) on the diagram of $e^{54}\text{Cr}$ versus $\Delta^{17}\text{O}$ (fig. S20). Using ^{53}Mn - ^{53}Cr chronometry (^{53}Mn decays to ^{53}Cr with a half-life of 3.7 My), the data place the accretion time of SM at 4566.57 ± 0.66 million years ago (Ma) (fig. S19) (28).

SM contains CM lithologies (from CM2.0 to 2.1, Fig. 2B) with varying histories of aqueous alteration and thermal metamorphism. The abundances of thermally labile elements, such as Se, Te, Zn, Sn, and Tl (Fig. 3 and table S7), known to be sensitive to open-system heating and volatilization, indicate that most of SM avoided metamorphic (on a million-year time scale) heating above 400° to 500°C (29). Like other CM chondrites, both lithologies of SM51 (Fig. 2B) and some of those in SM2 contain abundant carbonate grains and complete chondrule pseudomorphs embedded in a phyllosilicate-rich matrix (Fig. 2, B and C). Sample SM2-5 contains clasts of incompletely altered CM material, whose matrix consists largely of

submicrometer-sized olivine pseudomorphs after phyllosilicates and troilite (Fig. 2D and fig. S14). This and the complete lack of carbonates and tochilinite indicate that this particular clast has experienced thermal metamorphism to $\sim 500^\circ\text{C}$; all other metamorphosed CMs are finds (30).

Thermoluminescence (TL) measurements (31) of SM2-1d (fig. S35) show heating to $300^\circ \pm 20^\circ\text{C}$ within the last 0.2 My (fig. S36), and it has induced TL similar to low-metamorphic-grade CO and CV chondrites and unlike other CM chondrites. Raman spectra of macromolecular carbon (fig. S33) suggest that SM2-9 experienced only $153^\circ \pm 27^\circ\text{C}$, whereas a sample of SM12 experienced $268^\circ \pm 42^\circ\text{C}$ on a million-year time scale [using the method in (32)]. In a Raman G-band center-versus-width diagram (fig. S33), SM2-9 plots between CM2 and CO3 chondrites, whereas SM12 trends closer to polycrystalline C observed in CV3 chondrites.

There are various potential sources for the observed heating. The fast entry would have heated the meteorite surface to $>700^\circ\text{C}$ for up to 1.5 s, but it is unclear that this would offer enough time to alter the meteorite deeper inside. At 0.47 AU from the Sun, small tumbling asteroids warm up to $\sim 200^\circ\text{C}$, more at the surface, less inside. Alternatively, the recent heating could have been related to the impact that liberated the meteorite from the asteroid surface ≥ 0.05 Ma. Annealing

above 300°C may also have occurred on the parent body during the first 1 My after accretion, the most probable cause being the decay of live ^{26}Al .

SM, like Kaidun, shows large variations in O isotope compositions, reflecting the presence of diverse types of lithologies in both regolith breccias (Fig. 4). The aqueous alteration that produced carbonates and phyllosilicates caused variation in $\delta^{18}\text{O}$ along an approximately fixed $\Delta^{17}\text{O}$ ($=\delta^{17}\text{O} - 0.52 \times \delta^{18}\text{O}$) fractionation line (Fig. 4). This requires water-rock reaction involving flowing water as a result of a temperature gradient (33). Individual SM fragments plot at different positions along this line. The carbonates exhibit a very large range in $\delta^{18}\text{O}$ value [$+13$ per mil (‰) to $+39$ ‰] with $\Delta^{17}\text{O} = -1.9 \pm 1.5$ ‰ (fig. S21). A trend line regressed through the calcite data, $\delta^{17}\text{O} = -4.88 + 0.62 \times \delta^{18}\text{O}$, is nearly identical to that reported by (34) based on CM calcites.

SM has the lowest N/C ratio and $\delta^{15}\text{N}$, as compared to other CM2 chondrites (fig. S24), suggesting that it contains different N-bearing organic components. C and N isotopic compositions vary widely between fragments, with $\delta^{13}\text{C}$ and $\delta^{15}\text{N}$ values ranging between -13 ‰ and $+28.5$ ‰ and between -0.6 ‰ and $+16.7$ ‰, respectively. Bimodality in C release suggests two separate organic components, a volatile-rich and a volatile-poor component, the former being isotopically lighter and more cross-linked. Added to these components is CO_2 from

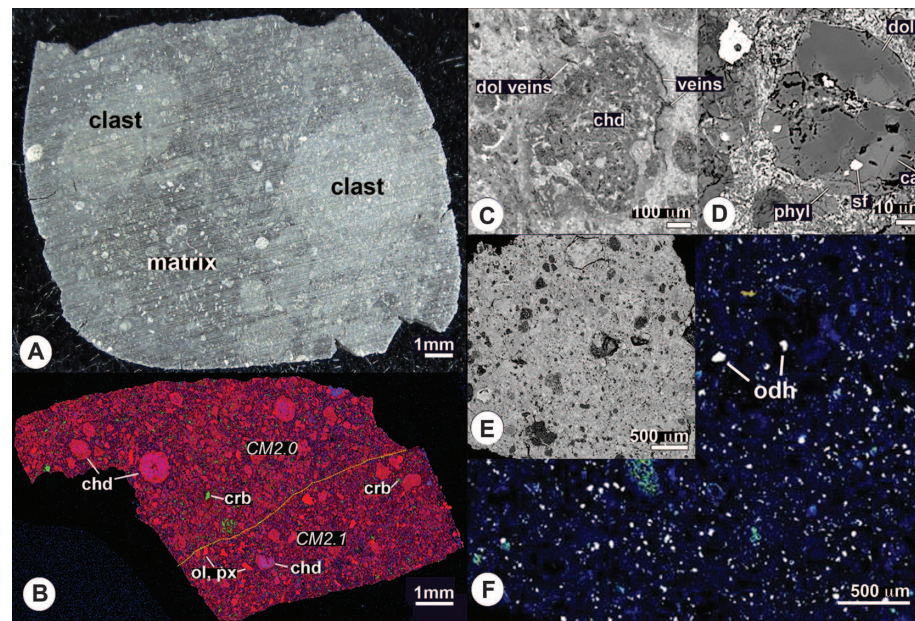


Fig. 2. (A) Slice of SM48 in visible light, showing light clasts in a dark matrix. (B) Combined elemental map in Mg (red), Ca (green), and Al (blue) Ka x-rays of the sample SM51-1 composed of two extensively aqueously altered CM2.0 and CM2.1 lithologies, with a sharp boundary (yellow dashed line). Both lithologies contain complete chondrule pseudomorphs (chd) embedded in a phyllosilicate-rich matrix and abundant carbonate grains (crb). In addition, the CM2.1 lithology contains rare olivine (ol) and pyroxene (px) grains of incompletely hydrated chondrules, amoeboid olivine aggregates (AOAs), and Ca-Al-rich inclusions (CAIs). (C) Backscattered electron (BSE) image of a chondrule pseudomorph (chd) composed mainly of phyllosilicates in SM51-1. The fine-grained rim around the chondrule is crosscut by veins of dolomite (dol), indicating in situ aqueous alteration. (D) BSE image of a carbonate grain composed of closely intergrown dolomite (dol) and calcite (cal) with inclusions of Fe,Ni-sulfide (sf) in SM51-1. The grain is rimmed by phyllosilicates (phyl). (E) BSE image of a small chip from sample SM2-5 with (F) corresponding Ca-S composite x-ray map (Ca, blue; S, green); all light spots in the map correspond to oldhamite (CaS; odh).

the decomposition of carbonates. The maximum in $\delta^{13}\text{C}$ of +65‰ is likely to be from the decomposition of calcite; $\delta^{13}\text{C}$, however, may be underestimated because of contamination by isotopically lighter components. About 2% of the total C combusts above 1000°C and reaches a $\delta^{13}\text{C}$ of +130‰. This is a well-known feature of carbonaceous chondrites, produced by the combustion of ^{13}C -enriched ($\delta^{13}\text{C} \sim +1400\text{‰}$) presolar silicon carbide grains (35). Apparently, SM contains a significant abundance of presolar grains.

Reactive compounds. The rapid pre-rain recovery of SM offers a rare glimpse of what reactive minerals and organic compounds may be present at the surface of asteroids. Unseen in other CM chondrites, abundant CaS grains were found in SM2-4 (Fig. 2E). A powder Laue pattern of these CaS grains proved to index as oldhamite (fig. S16). Oldhamite is quickly lost to moisture. The oldhamite grains are set within fine-grained comminuted matrix, containing also olivine, enstatite, Fe-Ni-Zn sulfides, Fe-Ni-Cr phosphides, and grains of reduced C, suggesting

admixture of a reduced component, possibly xenolithic enstatite chondrite material. It is interesting that the comminuted matrix of SM contains fragments of such a rare meteorite type, rather than the far more abundant ordinary chondrites, probably implying interactions between C- and E-type asteroids.

In comparison to SM2 (pre-rain), the SM12 sample (post-rain) shows strong alteration effects, demonstrating rapid reaction between terrestrial water and reactive S-bearing species. Fourier transform ion cyclotron resonance mass spectra (FTICR-MS) (36) of SM's methanol soluble fraction showed only hundreds of mass signals with comparatively low intensity, high aliphaticity, and, different from the thousands analyzed in other carbonaceous chondrites, many highly oxygenated species and polysulfur-rich compounds (fig. S31). Nuclear magnetic resonance (NMR) spectroscopy of the same extracts showed abundant highly branched, singly oxygenated aliphatics and a considerable diversity of unsaturated compounds in an intensity ratio of near 40:6:1 (fig. S32). Both NMR and FTICR-MS signals confirmed convergence with structures recently observed in rather highly thermally altered meteorites.

Water-soluble organic compounds and inorganic salts can be formed, mobilized, and altered by aqueous alteration and terrestrial weathering. Ion chromatography (37) of ≤ 20 mg of water-extracted pre-rain SM2 showed formate at 80 parts per million (ppm) and acetate at 700 ppm, whereas post-rain interior SM12 had trace formate and only ~ 100 ppm acetate. Detected inorganic anions in SM2 were sulfate (1300 ppm) and chloride (262 ppm), again with only trace amounts present in interior SM12. Murchison (presumably more aqueously altered) has a 20 times higher soluble sulfate abundance (26,000 ppm) (37) than SM2. Sodium was the dominant cation (on a per-mole basis) at ~ 1900 ppm in SM2, followed by calcium (2080 ppm) and magnesium (117 ppm).

Water extracts of SM2 were analyzed by gas chromatography mass spectrometry for ammonia and amines; amino, hydroxyl, and dicarboxylic acids (38); and, for amino acids, also by liquid chromatography with fluorimetric detection and time-of-flight mass spectrometry (39). The most abundant water-soluble compounds detected by these methods included glycine, β -alanine, γ -amino-*n*-butyric acid, and, in some analyses only, ϵ -amino-*n*-caproic acid (table S21). The highest total amino acid abundances were found in a fragment of SM12 at a depth of 9 to 12 mm from the crust. Exterior portions of SM2 and SM12 contained predominantly contaminant L-amino acids, and both samples were highly depleted in amino acids overall. The pH values of SM water extracts ranged from 8.90 to 9.65, which are higher than for other CMs.

Dichloromethane/methanol (9:1 volume to volume) extracts also showed variably low amounts of soluble hydrocarbon (fig. S34). SM2 contained naphthalene, methylnaphthalenes, dimethylnaphthalenes, anthracene/phenanthrene, a series of linear 15-C to 22-C alkanes, and very little S, usually abundant in comparable CM extracts (40). The SM12 extracts were instead dominated by cyclic

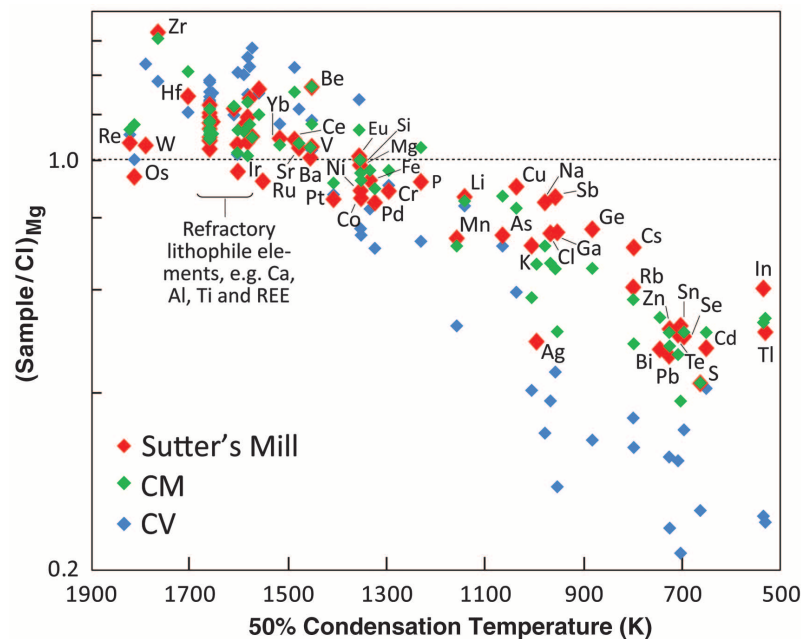


Fig. 3. Average elemental composition of SM meteorite (table S7) compared to averages for CI (Ivuna type), CM, and CV (Vigarano type) groups of carbonaceous chondrites. Data are normalized to CI and Mg and plotted against 50% condensation temperatures of the elements (41). Data sources: CI (41), CV (42), and CM (43).

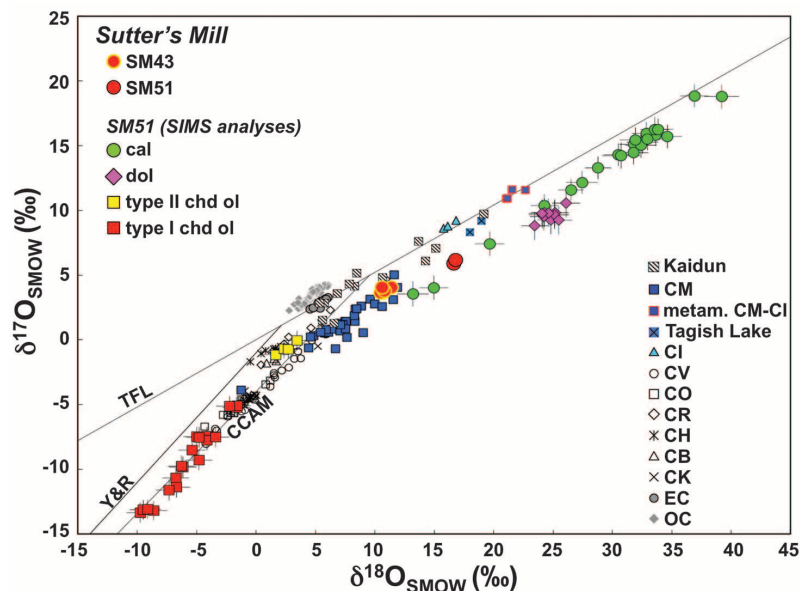


Fig. 4. SM43 and SM51 extend the currently known CM field in the three O-isotope diagram. Data for ordinary, enstatite, and carbonaceous chondrites and for Kaidun chondritic regolith breccia are from (43–47). TFL, terrestrial fractionation line; CCAM, carbonaceous chondrite anhydrous mineral line (44); Y&R, Young and Russell line (48). SMOW, standard mean ocean water.

octatomic S, dimethyltrisulfide, and dimethyl-tetrasulfide, the latter two for about a combined 1 mmole/g; the fragment contained in addition only naphthalene at 2 to 8 nmol/g, plus methyl-naphthalenes and biphenol in subnanomole amounts, but no alkanes or anthracene/phenanthrene.

The SM meteorite demonstrates that the complexity of C-class asteroid surfaces is greater than previously assumed. Rapid terrestrial alteration probably erases many vestiges of the internal and external processes on the asteroid that remain to be explored in spacecraft sample-return missions.

References and Notes

- M. Fries, J. Fries, *Meteorit. Planet. Sci.* **45**, 1476 (2010).
- P. Jenniskens, M. Zolensky, *Meteorit. Bull.* **46**, 1 (2012).
- Materials and methods are available as supplementary materials on Science Online.
- D. R. Christie, P. Campus, in *Infrasound Monitoring for Atmospheric Studies*, A. Le Pichon, E. Blanc, A. Hauchecorne, Eds. (Springer, Dordrecht, Netherlands, 2010), pp. 29–75.
- T. A. Ens, P. G. Brown, W. N. Edwards, E. A. Silber, *J. Atmos. Sol. Terr. Phys.* **80**, 208 (2012).
- E. A. Silber, A. Le Pichon, P. Brown, *Geophys. Res. Lett.* **38**, L12201 (2011).
- M. H. Shaddad *et al.*, *Meteorit. Planet. Sci.* **45**, 1557 (2010).
- K. D. Smith *et al.*, *Science* **305**, 1277 (2004).
- H. Haack *et al.*, *Meteorit. Planet. Sci.* **47**, 30 (2012).
- H. Haack *et al.*, in *Workshop on the First Solids in the Solar System*, 7 to 9 November 2011, Koloa, Kauai, Hawai'i (Lunar and Planetary Institute, Houston, TX, 2011), abstr. 9100.
- K. Nishiizumi, M. W. Caffee, *Meteorit. Planet. Sci.* **44** (Suppl.), 5358 (2009).
- U. Ott, *Rev. Mineral. Geochem.* **47**, 71 (2002).
- V. S. Heber *et al.*, *Geochim. Cosmochim. Acta* **73**, 7414 (2009).
- I. Leya, H.-J. Lange, S. Neumann, R. Wieler, R. Michel, *Meteorit. Planet. Sci.* **35**, 259 (2000).
- M. Gounelle *et al.*, in *The Solar System Beyond Neptune*, M. A. Barucci, H. Boehnhardt, D. P. Cruikshank, A. Morbidelli, Eds. (Univ. of Arizona Press, Tucson, AZ, 2008), pp. 525–541.
- D. Nesvorný *et al.*, *Astrophys. J.* **713**, 816 (2010).
- K. J. Walsh, M. Delbo, W. F. Bottke, paper presented at the 44th American Astronomical Society/Division of Planetary Sciences Meeting, Reno, NV, 14 to 19 October, 2012, abstr. 305.04.
- A. Tsuchiyama *et al.*, *Meteorit. Planet. Sci.* **44** (suppl.), 5189 (2009).
- O. Popova *et al.*, *Meteorit. Planet. Sci.* **46**, 1525 (2011).
- D. S. Ebel, M. L. Rivers, *Meteorit. Planet. Sci.* **42**, 1627 (2007).
- R. J. Macke, D. T. Britt, G. J. Consolmagno, *Meteorit. Planet. Sci.* **46**, 311 (2011).
- D. T. Britt *et al.*, paper presented at the Meteoritical Society Meeting, Cairns, Australia, 12 to 17 August 2012, abstr. 5250.
- M. S. Spiegel, R. C. Reedy, O. W. Lazareth, P. W. Levy, L. A. Slate, *J. Geophys. Res.* **91**, 483 (1986).
- P. Rochette *et al.*, *Meteorit. Planet. Sci.* **43**, 959 (2008).
- J. Gattacceca, P. Rochette, *Earth Planet. Sci. Lett.* **227**, 377 (2004).
- A. Bischoff, E. R. D. Scott, K. Metzler, C. A. Goodrich, in *Meteorites and the Early Solar System II*, D. S. Lauretta, H. Y. McSween Jr., Eds. (Univ. of Arizona Press, Tucson, AZ, 2006), pp. 679–712.
- R. J. Walker *et al.*, *Geochim. Cosmochim. Acta* **66**, 4187 (2002).
- Q.-Z. Yin, K. Yamashita, A. Yamakawa, R. Tanaka, B. Jacobsen, D. S. Ebel, I. D. Hutcheon, and E. Nakamura, *Lunar Planet. Sci.* **XL**, abstr. 2006 (2009).
- J. M. Friedrich, M.-S. Wang, M. E. Lipschutz, *Meteorit. Planet. Sci.* **37**, 677 (2002).
- E. Tonui, M. E. Zolensky, M. E. Lipschutz, *Proc. NIPR Symp. Antarct. Meteorites* **15**, 38 (2002).
- M. Haq, F. A. Hasan, D. W. G. Sears, C. B. Moore, C. F. Lewis, *Geochim. Cosmochim. Acta* **53**, 1435 (1989).
- G. D. Cody *et al.*, *Earth Planet. Sci. Lett.* **272**, 446 (2008).
- E. D. Young, R. D. Ash, P. England, D. Rumble 3rd, *Science* **286**, 1331 (1999).
- G. K. Benedix, L. A. Leshin, J. Farquhar, T. Jackson, M. H. Thieme, *Geochim. Cosmochim. Acta* **67**, 1577 (2003).
- M. M. Grady, *Astron. Geophys.* **50**, 4.21 (2009).
- P. Schmitt-Kopplin *et al.*, *Proc. Natl. Acad. Sci. U.S.A.* **107**, 2763 (2010).
- G. W. Cooper, W. M. Onwo, J. R. Cronin, *Geochim. Cosmochim. Acta* **56**, 4109 (1992).
- A. A. Monroe, S. Pizzarello, *Geochim. Cosmochim. Acta* **75**, 7585 (2011).
- D. P. Glavin, M. P. Callahan, J. P. Dworkin, J. E. Elsila, *Meteorit. Planet. Sci.* **45**, 1948 (2010).
- J. R. Cronin, S. Pizzarello, *Geochim. Cosmochim. Acta* **54**, 2859 (1990).
- K. Lodders, *Astrophys. J.* **591**, 1220 (2003).
- E. Jarosewich, R. S. Clarke Jr., J. N. Barrows, Eds., *Smithson. Contrib. Earth Sci.* **27**, 1 (1986).
- P. G. Brown *et al.*, *Science* **290**, 320 (2000).
- R. N. Clayton, T. K. Mayeda, *Geochim. Cosmochim. Acta* **63**, 2089 (1999).
- R. N. Clayton, T. K. Mayeda, J. N. Goswami, E. J. Olsen, *Geochim. Cosmochim. Acta* **55**, 2317 (1991).
- R. N. Clayton, T. K. Mayeda, A. E. Rubin, *J. Geophys. Res.* **89**, C245 (1984).
- G. J. MacPherson *et al.*, *Geochim. Cosmochim. Acta* **73**, 5493 (2009).
- E. D. Young, S. S. Russell, *Science* **282**, 452 (1998).

Acknowledgments: The rapid recovery of SM was made possible by the local communities of Lotus, Coloma, Pilot Hill, and Grass Valley, due in part to the support of property owners, community leaders, SETI Institute and University of California at Davis (UC Davis) student and volunteer searchers, and officials of the Marshall Gold Discovery State Historic Park and the State of California. We thank J. Howard, M. Carter, D. E. Macon, and L. Hofland of the NASA Ames Research Center; M. Hill of the American Museum of Natural History; and G. Baxter, N. Botto, D. Rowland, O. Kreylos, L. Kellogg, and D. Summer of UC Davis for assistance with reported measurements. The NASA Ames Research Center-coordinated search efforts were supported by the NASA Lunar Science Institute. E.D.Y., M.Z., A.M.D., D.S.E., J.E.E., K.N., K.C.W., M.H.T., M.W.C., Q.-Z.Y., and S.B.S. acknowledge the NASA Cosmochemistry Program for support of work reported here. In addition, the following NASA programs are sincerely acknowledged: Near Earth Object Observation Program (P.J.), Origins of Solar Systems (S.A.S.), Planetary Major Equipment (Q.-Z.Y.), Sample Return Laboratory Instruments and Data Analysis (M.W.C.), the NASA Astrobiology Institute (J.E.E., J.P.D., and D.P.G.), the Spring-8 Grant Program (M.Z.), and Post Doctoral Programs (A.S.B.). P.R.H. acknowledges support from the Tawani Foundation. R.O. acknowledges KAKENHI grants from the Japan Society for the Promotion of Science.

Supplementary Materials

www.sciencemag.org/cgi/content/full/338/6114/1583/DC1
Materials and Methods
Supplementary Text
Figs. S1 to S36
Tables S1 to S22
References (49–105)

9 July 2012; accepted 12 November 2012
10.1126/science.1227163

The Evolutionary Landscape of Alternative Splicing in Vertebrate Species

Nuno L. Barbosa-Morais,^{1,2} Manuel Irimia,^{1*} Qun Pan,^{1*} Hui Y. Xiong,^{3*} Serge Gueroussov,^{1,4*} Leo J. Lee,³ Valentina Slobodeniuc,¹ Claudia Kutter,⁵ Stephen Watt,⁵ Recep Çolak,^{1,6} TaeHyung Kim,^{1,7} Christine M. Misquitta-Ali,¹ Michael D. Wilson,^{4,5,7} Philip M. Kim,^{1,4,6} Duncan T. Odom,^{5,8} Brendan J. Frey,^{1,3} Benjamin J. Blencowe^{1,4,†}

How species with similar repertoires of protein-coding genes differ so markedly at the phenotypic level is poorly understood. By comparing organ transcriptomes from vertebrate species spanning ~350 million years of evolution, we observed significant differences in alternative splicing complexity between vertebrate lineages, with the highest complexity in primates. Within 6 million years, the splicing profiles of physiologically equivalent organs diverged such that they are more strongly related to the identity of a species than they are to organ type. Most vertebrate species-specific splicing patterns are cis-directed. However, a subset of pronounced splicing changes are predicted to remodel protein interactions involving trans-acting regulators. These events likely further contributed to the diversification of splicing and other transcriptomic changes that underlie phenotypic differences among vertebrate species.

Vertebrate species possess diverse phenotypic characteristics, yet they share similar repertoires of coding genes (*1*).

Evolutionary changes in transcriptomes underlie structural and regulatory differences associated with species-specific characteristics. For

example, species-dependent mRNA and non-coding RNA (ncRNA) expression patterns have been linked to mutational changes in cis- and trans-acting regulatory factors, as well as to phenotypic differences (2–5). However, because organ-dependent mRNA expression levels within individual species have been largely conserved during vertebrate evolution (6, 7), it seems unlikely that changes in gene expression (GE)

¹Banting and Best Department of Medical Research, Donnelly Centre, University of Toronto, Toronto, Ontario M5S 3E1, Canada. ²Instituto de Medicina Molecular, Faculdade de Medicina, Universidade de Lisboa, 1649-028 Lisboa, Portugal. ³Department of Electrical and Computer Engineering, University of Toronto, Toronto, Ontario M5S 3G4, Canada. ⁴Department of Molecular Genetics, University of Toronto, Toronto, Ontario M5S 1A8, Canada. ⁵University of Cambridge, CRUK-Cambridge Institute, Robinson Way, Cambridge CB2 0RE, UK. ⁶Department of Computer Science, University of Toronto, Toronto, Ontario M5S 2E4, Canada. ⁷Hospital for Sick Children, Toronto, Ontario M5G 1X8, Canada. ⁸Wellcome Trust Sanger Institute, Genome Campus, Hinxton, Cambridge CB10 1SA, UK.

*These authors contributed equally to this work.

†To whom correspondence should be addressed. E-mail: b.blencowe@utoronto.ca